

## DISCOVERY OF A RED QUASAR WITH RECURRENT ACTIVITY

S. NANDI<sup>1,2,3</sup>, R. ROY<sup>4</sup>, D.J. SAIKIA<sup>5,6</sup>, M. SINGH<sup>2</sup>, H.C. CHANDOLA<sup>3</sup>, M. BAES<sup>1</sup>, R. JOSHI<sup>2</sup>, G. GENTILE<sup>1,7</sup>, AND M. PATGIRI<sup>8</sup>

<sup>1</sup> Sterrenkundig Observatorium, Universiteit Gent, Krijgslaan 281 S9, B-9000 Gent, Belgium; [sumana1981@gmail.com](mailto:sumana1981@gmail.com), [Sumana.Nandi@UGent.be](mailto:Sumana.Nandi@UGent.be)

<sup>2</sup> Aryabhata Research Institute of Observational Sciences (ARIES), Manora Peak, Nainital, 263002, India

<sup>3</sup> Department of Physics, Kumaun University, Nainital 263001, India

<sup>4</sup> Institut d'Astrophysique et de Géophysique, Université de Liège, Allée du 6 Août 17, Bât B5c, B-4000 Liège, Belgium

<sup>5</sup> National Centre for Radio Astrophysics, TIFR, Pune University Campus, Post Bag 3, Pune 411007, India

<sup>6</sup> Cotton College State University, Panbazar, Guwahati 781001, India

<sup>7</sup> Department of Physics and Astrophysics, Vrije Universiteit Brussel, Pleinlaan 2, B-1050 Brussels, Belgium

<sup>8</sup> Cotton College, Panbazar, Guwahati 781001, India

Received 2014 March 26; accepted 2014 May 6; published 2014 June 10

### ABSTRACT

We report a new double–double radio quasar (DDRQ) J0746+4526 which exhibits two cycles of episodic activity. From radio continuum observations at 607 MHz using the Giant Metrewave Radio Telescope and 1400 MHz from the Faint Images of the Radio Sky at Twenty-cm survey we confirm its episodic nature. We examine the Sloan Digital Sky Survey (SDSS) optical spectrum and estimate the black hole mass to be  $(8.2 \pm 0.3) \times 10^7 M_{\odot}$  from its observed Mg II emission line, and the Eddington ratio to be 0.03. The black hole mass is significantly smaller than for the other reported DDRQ, J0935+0204, while the Eddington ratios are comparable. The SDSS spectrum is significantly red-continuum-dominated, suggesting that it is highly obscured with  $E(B - V)_{\text{host}} = 0.70 \pm 0.16$  mag. This high obscuration further indicates the existence of a large quantity of dust and gas along the line of sight, which may have a key role in triggering the recurrent jet activity in such objects.

*Key words:* galaxies: active – galaxies: individual (J0746+4526) – galaxies: nuclei – radio continuum: galaxies

### 1. INTRODUCTION

It is well established that many active galactic nuclei (AGNs) may go through two or more cycles of episodic activity (Subrahmanyan et al. 1996; Lara et al. 1999). This is most clearly seen in extended radio galaxies and quasars, where there may be a new pair of radio lobes with well-defined hot spots closer to the nucleus, in addition to the more distant and diffuse, extended lobes from an earlier cycle of activity. The sources with a second pair of lobes have been classified as double–double radio galaxies (DDRGs) by Schoenmakers et al. (2000), and their properties have been summarized by Saikia & Jamrozy (2009). Although most DDRGs were believed initially to be giant radio sources with sizes of over about a megaparsec, a significant number of smaller sized DDRGs have also been identified in recent years (Nandi & Saikia 2012). In most of the cases double–double radio sources are associated with galaxies but it is also possible for a quasar to appear as a radio source exhibiting episodic activity. However, the number of double–double radio quasars (DDRQs) reported so far is very limited (Jamrozy et al. 2009), and it is important to identify more of these to make the physical scenario of this class statistically robust. J0935+0204 (4C02.27), located at the redshift of 0.65, is one promising DDRQ reported by Jamrozy et al. (2009). It has a blue-continuum-dominated spectrum, with projected linear sizes of 70 and 470 kpc for the inner and outer radio lobes, respectively. Although in optical surveys the widely popular “Color selection technique” already identified a large number of quasars having blue-continuum-dominated spectra, recent studies found huge dust accumulation around many new similar objects, which make them optically redder (Urrutia et al. 2009). These red quasars are usually believed to represent an early stage of AGN evolution, which is going through a merging process, leading to the presence of large amounts of gas and dust (Glikman et al. 2012).

Based on only radio structural information Proctor (2011) classified 811,117 radio sources or entries from the FIRST (Faint

Images of the Radio Sky at Twenty-cm; Becker et al. 1995) survey into different categories and listed 242 sources as candidate DDRGs. Further detailed investigation of these sources along with optical data from SDSS<sup>9</sup> and DSS<sup>10</sup> catalogues showed only 23 of these sources to be promising examples of DDRGs (Nandi & Saikia 2012). In this paper we report the discovery of a red DDRQ J0746+4526 and discuss some of its properties and also compare these with the properties of J0935+0204. J0746+4526 which was classified from this survey to show evidence of episodic activity, has been identified as a quasar at a redshift  $0.55021 \pm 0.00005$  in Data Release 9 of Sloan Digital Sky Survey (SDSS). However, to confirm that neither of the inner compact components is a compact flat-spectrum nuclear component, which appears to be the case in many of the candidate DDRGs identified by Proctor (2011), we observed this source with the Giant Metrewave Radio Telescope (GMRT) at 607 MHz.

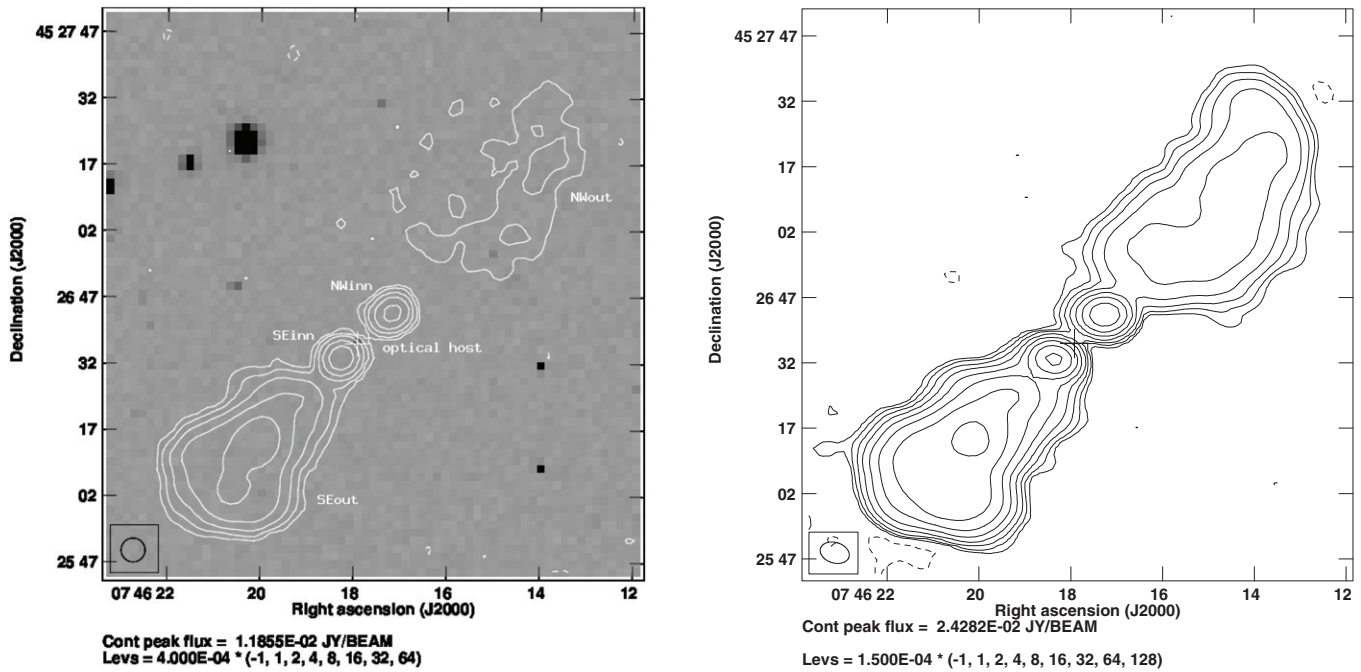
We present the results of these observations of J0746+4526, which confirm the source to be a DDRQ, and then discuss the nature of the source using the results of our observations and archival FIRST data at radio wavelengths and optical data from SDSS. The observations and radio data reduction procedures are described in Section 2. In Section 3, we present the observational results. The black hole (BH) mass determination is described in Section 4. The discussion and concluding remarks are summarized in Section 5.

### 2. OBSERVATIONS AND DATA REDUCTION

We made the 607 MHz GMRT observations of J0746+4526 on 2012 November 29. The observations were made in the standard manner, with each observation of the target source interspersed with observations of the phase calibrator, J0713+438.

<sup>9</sup> <http://www.sdss3.org/dr9/>

<sup>10</sup> <http://archive.eso.org/dss/dss>



**Figure 1.** Left panel: the FIRST image at 1400 MHz overlaid on the optical field from the SDSS. The optical host and both the outer and inner doubles (NW<sub>out</sub>, SE<sub>out</sub>, NW<sub>inn</sub>, SE<sub>inn</sub>) are marked in the same image. Right panel: the GMRT image at 607 MHz. The + sign represents the host position.

After each of several 20 minute exposures of the target source, the phase calibrator was observed for 5 minutes. The total observing time on the target source is about 190 minutes. We observed 3C 147 for flux density as well as bandpass calibration. All flux densities are on the Baars et al. (1977) scale using the latest Very Large Array values. These data were edited and reduced with the NRAO AIPS package. Several rounds of self-calibration were done to produce the best possible images.

### 3. OBSERVATIONAL RESULTS

#### 3.1. Radio Data

The full-resolution radio image obtained from FIRST at 1400 MHz and that from GMRT at 607 MHz are presented in Figure 1. The diffuse northwestern outer lobe is imaged better at the lower-frequency GMRT image. The observational parameters and the flux densities estimated from these images are presented in Table 1. The spectral indices  $\alpha$  (defined as  $S_\nu \propto \nu^{-\alpha}$ ) between 607 and 1400 MHz for the NW<sub>out</sub> (northwest outer) and SE<sub>out</sub> (southeast outer) lobes are  $1.62 \pm 0.13$  and  $1.18 \pm 0.13$ , respectively. The corresponding two-point spectral indices for the NW<sub>inn</sub> (northwest inner) and SE<sub>inn</sub> (southeast inner) components are  $0.93 \pm 0.13$  and  $1.00 \pm 0.13$ , respectively.

The outer doubles of this DDRQ appear reasonably well aligned with the inner ones and are collinear with the parent optical host galaxy. The radio core has not been detected at either frequency. The projected linear size between the outer radio lobes is  $\sim 630$  kpc whereas the separation between two inner lobes is  $\sim 95$  kpc.<sup>11</sup> The outer lobes are highly asymmetric in intensity and do not show any evidence of hot spots at the ends of the lobes. The flux density ratio of the outer components at 1400 MHz is 4.96, while for the inner lobes it is 1.18, with the southeastern component being brighter in both cases. The log

<sup>11</sup> We have assumed a universe with  $H_0 = 71 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ,  $\Omega_m = 0.27$ , and  $\Omega_{\text{vac}} = 0.73$

**Table 1**  
The Observational Parameters Obtained from Radio Data<sup>a</sup>

Freq. (MHz)	Beam Size (")		P.A. (°)	rms (mJy/b)	Comp.	$S_p$ (mJy/b)	$S_f$ (mJy)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)
607	6.83	4.71	73	0.05	NW <sub>out</sub>	3	93
					NW <sub>in</sub>	17	24
					SE <sub>in</sub>	24	30
					SE <sub>out</sub>	24	319
1400	5.40	5.40	0	0.14	NW <sub>out</sub>	1	24
					NW <sub>in</sub>	8	11
					SE <sub>in</sub>	12	13
					SE <sub>out</sub>	10	119

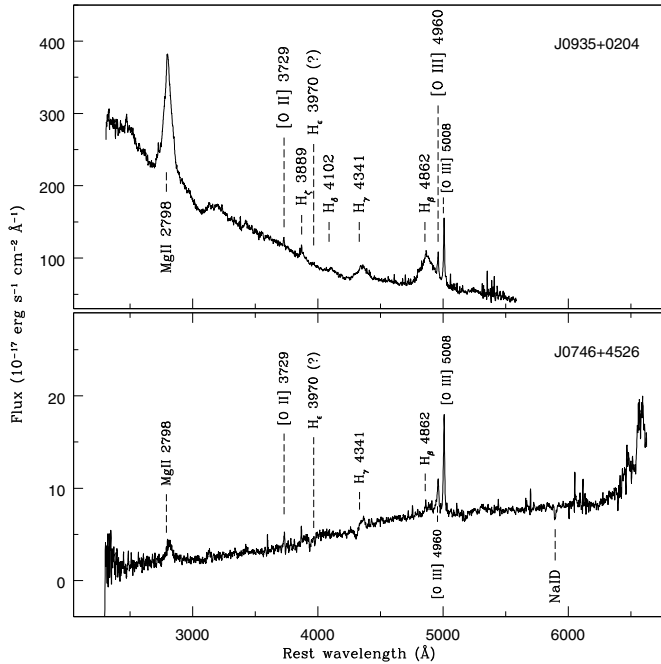
**Note.** <sup>a</sup> Arrangement of the table is as follows. Column 1: frequencies of observations; Columns 2 and 3: the angular sizes of the major and minor axes of the restoring beam; Column 4: the position angle; Column 5: the rms noise; Column 6: component designation; Column 7: the peak flux density; Column 8: the total flux density of each component. The error in the flux density is approximately 7% at 607 MHz and 5% at 1400 MHz.

luminosity at an emitted frequency of 1400 MHz for the inner and outer doubles are 25.41 and 26.17  $\text{W Hz}^{-1}$ , respectively (Nandi & Saikia 2012).

#### 3.2. Optical Data

Optical analysis is based on archival spectra obtained from the SDSS DR9 catalog for both DDRQs J0746+4526 and J0935+0204. The Galactic reddening  $E(B - V)$ , adopted from the NASA Extragalactic Database (NED), toward J0746+4526 and J0935+0204 are, respectively, 0.056 and 0.042 mag.

The Galactic extinction and redshift-corrected optical spectrum of J0746+4526 is presented in Figure 2, along with the spectrum of J0935+0204. The spectrum of J0935+0204 is similar to the typical spectra of normal quasars, while the spectrum of J0746+4526 is typical of red quasars. For both quasars, the



**Figure 2.** SDSS spectra of J0935+0204 (upper panel) and J0746+4526 (lower panel). For both spectra, the wavelengths were corrected for the redshifts of the galaxies, while the fluxes were measured after performing the corrections due to corresponding redshifts and galactic extinctions. Several spectral lines have been marked in both spectra.

[O III] lines at  $\lambda 4960 \text{ \AA}$  and  $\lambda 5008 \text{ \AA}$  associated with the narrow-line-region (NLR) clouds are quite prominent, while there is a striking difference in the strength of the hydrogen Balmer lines which are weaker in J0746+4526. It should be noted that the spectra of both quasars are of moderate signal-to-noise ratio. For J0746+4526 and J0935+0204 it is  $\sim 13.6$  and  $\sim 26.5$ , respectively, measured around the middle portion of the SDSS spectral range coverage.

We estimated the reddening for the object J0746+4526 in two different ways. First we use the empirical relation, developed from the study of spectra of nearby supernovae (e.g., Barbon et al. 1990; Turatto et al. 2003; Roy et al. 2011; Poznanski et al. 2012), between the width of the Na I D absorption dip and the host reddening. This has been widely used to determine the host galaxy extinction directly from the observed spectrum. The presence of the Na I D lines ( $\lambda\lambda 5890, 5896 \text{ \AA}$ ) in the spectrum of J0746+4526, for which we obtain an equivalent width of  $3.45 \text{ \AA}$ , enables us to estimate the plausible host extinction directly from the observed spectrum. Using the relation prescribed by Turatto et al. (2003), this corresponds to a host reddening of approximately 0.54 mag. On the other hand, according to the relation proposed by Barbon et al. (1990), the reddening of the host is  $\sim 0.86$  mag. We therefore assume that the mean host reddening along the line of sight is  $E(B - V)_{\text{host}} = 0.70 \pm 0.16$  mag.

We have alternatively estimated the amount of foreground extinction by comparing the de-reddened spectrum of J0746+4526 with an appropriate template (composite spectrum obtained from Brotherton et al. 2001<sup>12</sup>). To do so we have adopted the Small Magellanic Cloud extinction curve, and we have de-reddened our observed spectrum by considering dust distributed in front of the QSO in a uniform screen geometry.

We have searched for the best value of the color excess that minimizes the differences between the template and our QSO spectrum. We find that the host reddening is  $\sim 0.79$  mag. This result is consistent with the result obtained from the Na I D absorption dip. This is also consistent with the typical reddening ( $\sim 0.1$ – $1.5$  mag) determined for the red quasars (Urrutia et al. 2009).

The widths of several spectral lines which are presumed to be generated near the central region of the AGN (like the Mg II  $\lambda 2800 \text{ \AA}$  and H Balmer lines) are widely different for two interesting objects. For J0746+4526, the FWHM of Mg II  $\lambda 2800 \text{ \AA}$  is  $\sim 35.22 \text{ \AA}$ , whereas for J0935+0204 it is about  $61.33 \text{ \AA}$  (Kozieł-Wierzbowska et al. 2012). These demonstrate that the velocity dispersions of several broad-line-region (BLR) clouds are much higher for J0935+0204 in comparison to that for J0746+4526 and further indicate that the former galaxy hosts a higher mass central BH than the latter one.

#### 4. ESTIMATION OF THE BLACK HOLE MASS AND EDDINGTON RATIO

“Reverberation Mapping” (RM) or “Echo Mapping” has proven to be a viable technique to measure the location of the line emitting clouds and the BH mass (Blandford & McKee 1982; Peterson 1993). A major drawback with RM is that for a single system it requires a long-term monitoring program to measure the time gap between the continuum and the broad line variability. To measure the BH mass of this DDRQ we follow the indirect technique of the virial single-epoch method, which is an approximation of the RM method (Kaspi et al. 2005; Chand et al. 2010; Joshi et al. 2011; Kuźmicz & Jamrozny 2012). It exploits the empirical power-law correlation between the size of the BLR and the AGN continuum luminosity ( $R_{\text{BLR}} \propto L^\beta$ ; with  $\beta \sim 0.5$ ), as expected from photoionization model predictions (Bentz et al. 2006).

The best calibration values for the single epoch mass measurement scaling relations are available for the H $\beta$  line, as it is the basis for the majority of RM program. However, the similar  $R_{\text{BLR}}-L$  relationship is also frequently employed to measure the BH mass using the C IV and Mg II lines. The redshifts of these DDRQs are such that only the Mg II  $\lambda\lambda 2798, 2803$  and H $\beta$   $\lambda 4861$  lines fall in the wavelength range of the observed spectra. We note that the H $\beta$  and Mg II emission features are prominent in the spectrum of J0935+0204, whereas for J0746+4526 the Mg II line is clearly seen (although weaker) and the H $\beta$  emission line is barely seen (see Figure 2). Therefore, to determine the BH mass of J0746+4526, we have used the Mg II emission line, whereas for J0935+0204, Kuźmicz & Jamrozny (2012) used both the Mg II and H $\beta$  emission lines.

According to Vestergaard & Osmer (2009) (followed by Kuźmicz & Jamrozny 2012), the FWHM of the Mg II emission line of a quasar spectrum is related to its BH mass through the following scaling formula:

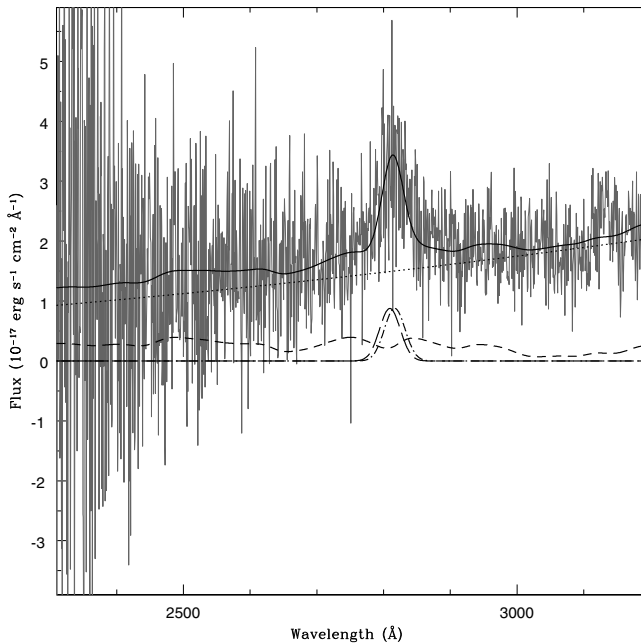
$$\frac{M_{\text{BH}}}{M_{\odot}} = 7.24 \times 10^6 \left[ \frac{\lambda L_{\lambda}(3000 \text{ \AA})}{10^{44} \text{ erg s}^{-1}} \right]^{0.5} \left[ \frac{\text{FWHM}_{\text{Mg II}}}{1000 \text{ km s}^{-1}} \right]^2, \quad (1)$$

where  $L_{\lambda}(3000 \text{ \AA})$  is the monochromatic luminosity at  $3000 \text{ \AA}$ .

The Mg II feature of the J0746+4526 spectrum has been fitted with multiple Gaussian functions along with a power-law continuum and underlying UV Fe II features have been obtained from the templates of nearby QSOs (Tsuzuki et al. 2006).

<sup>12</sup> <http://sundog.stsci.edu/first/QSOComposites/>.





**Figure 3.** Best fit to the Mg II emission line of the SDSS spectrum of J0746+4526. The observed spectrum is presented in gray. The continuum fit is presented by the dotted line. The long-dashed and dot-dashed lines represent the two broad emission features and the dashed line marks the underlying Fe II emission. The modeled spectrum is presented as a continuous line.

In Gaussian profile fits, our initial guess consists of two Gaussian components for each line of the Mg II doublet. However, if one of the components becomes statistically insignificant the procedure automatically drops it during the fit (Chand et al. 2010). In our two-component Gaussian profile fits, the width of each component (narrow/broad) of Mg II 2796 Å was tied to the respective components of the Mg II 2803 Å line. Here, the broad and narrow components represent the Mg II emission associated with the BLR and NLR clouds, respectively. The width of the broad components are set to be greater than  $2000 \text{ km s}^{-1}$  whereas those for the narrow components are constrained to be less than  $1000 \text{ km s}^{-1}$ . In addition, we have constrained the width of the broader Mg II component to be the same as the width of the UV Fe II emission line in this region.

The fitting has been done in an iterative way, minimizing the reduced  $\chi^2$  to  $\approx 1$ . Our final spectral fitting in the UV region is shown in Figure 3. The absence of the narrow Mg II features in this fitting process probably indicates that no (or very negligible) Mg II emissions are produced by the NLR clouds associated with J0746+4526. The contribution of the underlying Fe II features is also very small. The FWHM of the broad components, obtained from the fitting, is  $35.22 \pm 0.1 \text{ \AA}$ . This corresponds to a BH mass of  $(8.2 \pm 0.3) \times 10^7 M_{\odot}$ , where the error has been estimated after propagating the errors associated with the flux estimation at 3000 Å and FWHM estimation of the Mg II line in quadrature. This value is about one order smaller than the BH mass estimated for J0935+0204 ( $(13.22 \pm 0.75) \times 10^8 M_{\odot}$  from Mg II fit; Kuźmicz & Jamrozy 2012).

The Eddington ratio was computed from the relation  $\ell = L_{\text{bol}}/L_E$ , where  $L_{\text{bol}}$  is the bolometric luminosity and  $L_E$  is the Eddington luminosity (Metcalf & Magliocchetti 2006; Dietrich et al. 2009; Kuźmicz & Jamrozy 2012). For J0746+4526, the computed values of  $L_{\text{bol}}$  and  $L_E$  are, respec-

tively,  $\sim 3.7 \times 10^{44} \text{ erg s}^{-1}$  and  $\sim 1.2 \times 10^{46} \text{ erg s}^{-1}$ . This implies, for J0746+4526, that the Eddington ratio  $\sim 0.03$ , which is consistent with that of J0935+0204 ( $\sim 0.06$  calculated from Mg II profiles; Kuźmicz & Jamrozy 2012).

## 5. DISCUSSION AND CONCLUSIONS

In this paper we presented a radio and optical spectroscopic study of the DDRQ J0746+4526. Both GMRT and FIRST images of this object show that there were two episodes of activity. The diffuse emission of the outer double is well beyond both hot spots of the inner lobes. The steeper spectral indices of the outer lobes suggest that these are due to older emission. The lobes of the outer double are significantly more asymmetric in flux density than the inner ones. The core emission is not detected at either 1400 or 607 MHz. This is in contrast with the DDRQ J0935+0204, where a prominent core has been detected at both frequencies. Although higher resolution observations at higher frequencies would be required to estimate the core flux density reliably, non-detection of a core in J0746+4526 is consistent with a larger angle of inclination of this source to the line of sight (Saikia & Kulkarni 1994).

Although the projected linear size of J0746+4526 is larger than that of J0935+0204, both are FR II type, lobe-dominated DDRQs and highly asymmetric in flux density. J0935+0204 has a bright hot spot in the southwestern outer lobe. This feature indicates that, for this source, the timescale of episodic activity is less than a few million years (Saikia & Jamrozy 2009; Jamrozy et al. 2009). On the other hand, both outer lobes of J0746+4526 do not show any hot spots, and a more detailed multi-frequency study would be required to estimate its ages.

J0746+4526 and J0935+0204 are two quasars where recurrent jet activity has been observed. The optical spectrum of J0935+0204 shows a blue-continuum-dominated spectrum, similar to most quasars, whereas that of J0746+4526 shows evidence of more extinction and obscuration, similar to that of red quasars. Significant populations of obscured red quasars have been reported in recent studies (e.g., Urrutia et al. 2009; Glikman et al. 2012). These findings demonstrate that some different mechanism rather than the orientation of the torus is possibly responsible for large obscuration. In particular, the large extinction in red quasars is probably related to significant dust extinction in the host galaxies. The fact that out of the two known DDRQs one is a red quasar is an interesting result. As large dust extinction is often related to massive starbursts and galaxy merging, it suggests that galaxy merging can trigger not only powerful AGN activity (Heckman et al. 1986; Wilson & Colbert 1995; Sánchez & González-Serrano 2002) but also the interruption and restarting of jet formation. In other words, this result plausibly supports the merger-driven scenario for DDRG formation (Liu et al. 2003). In order to test whether this mechanism is the main scenario for DDRG formation, a larger sample of DDRGs needs to be investigated and compared with the global AGN population. The present paper, which presents the first discovery of a red quasar with recurrent jet activity, is only the first step toward a larger investigation.

The BH masses of giant and small-sized radio quasars, measured from Mg II emission lines, typically fall in the range  $1.6 \times 10^8 M_{\odot} < M_{\text{BH}} < 12.2 \times 10^8 M_{\odot}$  and  $1.0 \times 10^8 M_{\odot} < M_{\text{BH}} < 20.3 \times 10^8 M_{\odot}$ , respectively (Kuźmicz & Jamrozy 2012). For J0746+4526, the BH mass estimated using the Mg II

line is smaller than that of J0935+0204. However, their accretion rates are comparable. Although Kuźmicz & Jamrozy (2012) suggested that larger mass BHs are associated with smaller accretion rates, this does not persist for all types of quasars (e.g., J2335–0927, J1623+3419, J1433+3209; Kuźmicz & Jamrozy 2012). Hence, it appears that the BH mass and accretion rates of these DDRQs are quite similar to other quasars.

The discovery of a new DDRQ underlines the importance of identifying more of such objects to study their properties at different wavelengths to estimate timescales of episodic activity, and the properties of host galaxies and their environments to understand the triggering mechanisms for jet activity.

We are grateful to Hum Chand, Jean Surdej, and Jacopo Fritz for several useful discussions. We thank also the anonymous referee for her/his valuable comments. S.N. and R.R. are funded by the BELSPO grant to conduct their scientific researches. H.C.C. is grateful to the UGC (New Delhi) for financial assistance for the research project. The GMRT is a national facility operated by the National Centre for Radio Astrophysics of the Tata Institute of Fundamental Research. The National Radio Astronomy Observatory is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc. This research made use of the NASA/IPAC extragalactic database (NED) which is operated by the Jet Propulsion Laboratory, Caltech, under contract with the National Aeronautics and Space Administration. Funding for the SDSS and SDSS-II has been provided by the Alfred P. Sloan Foundation, the Participating Institutions, the National Science Foundation, the U.S. Department of Energy, the National Aeronautics and Space Administration, the Japanese Monbukagakusho, the Max Planck Society, and the Higher Education Funding Council for England.

## REFERENCES

- Baars, J. W. M., Genzel, R., Pauliny-Toth, I. I. K., & Witzel, A. 1977, *A&A*, **61**, 99
- Barbon, R., Benetti, S., Rosino, L., Cappellaro, E., & Turatto, M. 1990, *A&A*, **237**, 79
- Becker, R. H., White, R. L., & Helfand, D. J. 1995, *ApJ*, **450**, 559
- Bentz, M. C., Peterson, B. M., Pogge, R. W., Vestergaard, M., & Onken, C. A. 2006, *ApJ*, **644**, 133
- Blandford, R. D., & McKee, C. F. 1982, *ApJ*, **255**, 419
- Brotherton, M. S., Tran, H. D., Becker, R. H., et al. 2001, *ApJ*, **546**, 775
- Chand, H., Wiita, P. J., & Gupta, A. C. 2010, *MNRAS*, **402**, 1059
- Dietrich, M., Mathur, S., Grupe, D., & Komossa, S. 2009, *ApJ*, **696**, 1998
- Glikman, E., Urrutia, T., Lacy, M., et al. 2012, *ApJ*, **757**, 51
- Heckman, T. M., Smith, E. P., Baum, S. A., et al. 1986, *ApJ*, **311**, 526
- Jamrozy, M., Saikia, D. J., & Konar, C. 2009, *MNRAS*, **399**, L141
- Joshi, R., Chand, H., Gupta, A. C., & Wiita, P. J. 2011, *MNRAS*, **412**, 2717
- Kaspi, S., Maoz, D., Netzer, H., et al. 2005, *ApJ*, **629**, 61
- Koziel-Wierzbowska, D., Jamrozy, M., Zola, S., Stachowski, G., & Kuźmicz, A. 2012, *MNRAS*, **422**, 1546
- Kuźmicz, A., & Jamrozy, M. 2012, *MNRAS*, **426**, 851
- Lara, L., Márquez, I., Cotton, W. D., et al. 1999, *A&A*, **348**, 699
- Liu, F. K., Wu, X.-B., & Cao, S. L. 2003, *MNRAS*, **340**, 411
- Metcalf, R. B., & Magliocchetti, M. 2006, *MNRAS*, **365**, 101
- Nandi, S., & Saikia, D. J. 2012, *BASI*, **40**, 121
- Peterson, B. M. 1993, *PASP*, **105**, 247
- Poznanski, D., Prochaska, J. X., & Bloom, J. S. 2012, *MNRAS*, **426**, 1465
- Proctor, D. D. 2011, *ApJS*, **194**, 31
- Roy, R., Kumar, B., Moskvitin, A. S., et al. 2011, *MNRAS*, **414**, 167
- Saikia, D. J., & Jamrozy, M. 2009, *BASI*, **37**, 63
- Saikia, D. J., & Kulkarni, V. K. 1994, *MNRAS*, **270**, 897
- Sánchez, S. F., & González-Serrano, J. I. 2002, *A&A*, **396**, 773
- Schoenmakers, A. P., de Bruyn, A. G., Röttgering, H. J. A., van der Laan, H., & Kaiser, C. R. 2000, *MNRAS*, **315**, 371
- Subrahmanyan, R., Saripalli, L., & Hunstead, R. W. 1996, *MNRAS*, **279**, 257
- Tsuzuki, Y., Kawara, K., Yoshii, Y., et al. 2006, *ApJ*, **650**, 57
- Turatto, M., Benetti, S., & Cappellaro, E. 2003, in *Proc. ESO/MPA/MPE, From Twilight to Highlight: The Physics of Supernovae*, ed. W. Hillebrandt & B. Leibundgut (Berlin: Springer), **200**
- Urrutia, T., Becker, R. H., White, R. L., et al. 2009, *ApJ*, **698**, 1095
- Vestergaard, M., & Osmer, P. S. 2009, *ApJ*, **699**, 800
- Wilson, A. S., & Colbert, E. J. M. 1995, *ApJ*, **438**, 62