

***Herschel* photometric observations of the nearby low metallicity irregular galaxy NGC 6822^{★,★★}**

M. Galametz¹, S. C. Madden¹, F. Galliano¹, S. Hony¹, M. Sauvage¹, M. Pohlen², G. J. Bendo³, R. Auld², M. Baes⁴, M. J. Barlow⁵, J. J. Bock⁶, A. Boselli⁷, M. Bradford⁶, V. Buat⁷, N. Castro-Rodríguez⁸, P. Chanial¹, S. Charlot⁹, L. Ciesla⁷, D. L. Clements³, A. Cooray¹⁰, D. Cormier¹, L. Cortese², J. I. Davies², E. Dwek¹¹, S. A. Eales², D. Elbaz¹, W. K. Gear², J. Glenn¹², H. L. Gomez², M. Griffin², K. G. Isaak¹³, L. R. Levenson⁶, N. Lu⁶, B. O'Halloran³, K. Okumura¹, S. Oliver¹⁴, M. J. Page¹⁵, P. Panuzzo¹, A. Papageorgiou², T. J. Parkin¹⁶, I. Pérez-Fournon⁸, N. Rangwala¹², E. E. Rigby¹⁷, H. Roussel⁹, A. Rykala², N. Sacchi¹⁸, B. Schulz¹⁹, M. R. P. Schirm¹⁶, M. W. L. Smith², L. Spinoglio¹⁸, J. A. Stevens²⁰, S. Sundar⁹, M. Symeonidis¹⁵, M. Trichas³, M. Vaccari²¹, L. Vigroux⁹, C. D. Wilson¹⁶, H. Wozniak²², G. S. Wright²³, and W. W. Zeilinger²⁴

(Affiliations are available in the online edition)

Received 31 March 2010 / Accepted 21 April 2010

ABSTRACT

We present the first *Herschel* PACS and SPIRE images of the low-metallicity galaxy NGC 6822 observed from 70 to 500 μm and clearly resolve the H II regions with PACS and SPIRE. We find that the ratio 250/500 is dependent on the 24 μm surface brightness in NGC 6822, which would locally link the heating processes of the coldest phases of dust in the ISM to the star formation activity. We model the SEDs of some regions H II regions and less active regions across the galaxy and find that the SEDs of H II regions show warmer ranges of dust temperatures. We derive very high dust masses when graphite is used in our model to describe carbon dust. Using amorphous carbon, instead, requires less dust mass to account for submm emission due to its lower emissivity properties. This indicates that SED models including *Herschel* constraints may require different dust properties than commonly used. The global G/D of NGC 6822 is finally estimated to be 186, using amorphous carbon.

Key words. galaxies: ISM – galaxies: dwarf – galaxies: photometry

1. Introduction

The absorption of stellar radiation and its reemission by dust at infrared (IR) wavelengths is a fundamental process controlling the heating and cooling of the interstellar medium (ISM). The *IRAS*, *ISO* and *Spitzer* IR space telescopes launched the studies of the physics and chemistry of dust and gas, revealing their roles in the matter cycle and thermal balance in galaxies. The mid-infrared (MIR) to far-infrared (FIR) wavelength windows provide the necessary observational constraints on the spectral energy distribution (SED) modelling of galaxies from which properties of the polycyclic aromatic hydrocarbons (PAHs) and the warm (>30 K) and hot dust can be determined. Now, the *Herschel* Space Observatory (Pilbratt et al. 2010) is probing the submillimeter regime, where the coldest phases (<30 K) of dust can be revealed.

Dwarf galaxies of the Local Group are nearby laboratories to spatially study the life cycle of the different dust components and the metal enrichment of low-metallicity ISM. NGC 6822 is our closest (490 kpc, Mateo 1998) metal-poor galaxy neighbour (~30% Z_{\odot}) beyond the Magellanic Clouds and possesses isolated star forming (SF) regions at different evolutionary stages.

The galaxy is a perfect candidate to study the feedback of the star formation on the low-metallicity ISM by analysing the spatial variations of its SEDs. NGC 6822 also possesses an intriguing rotating H I disk of $1.34 \times 10^8 M_{\odot}$, that extends far beyond the optical disk (Mateo 1998) and has one of the largest HI holes ever observed in a dwarf galaxy (de Blok & Walter 2000).

NGC 6822 was observed in 2009 October as part of the science demonstration observations for the Dwarf Galaxy Survey (PI: S. Madden), with the instruments PACS and SPIRE at 70, 100, 160 and 250, 350 and 500 μm respectively. At SPIRE 500 μm (36''), we can resolve ISM structures of ~85 pc, spatially sufficient to accurately probe the distribution of dust temperature throughout the galaxy, especially dust in its coldest phases.

2. Observations and data reduction

PACS (Poglitsch et al. 2010) observations were performed in cross scan map mode at 70, 100 and 160 μm . The observations cover a region of 18' \times 18' around the starforming complexes of the galaxy also mapped with *Spitzer* (Cannon et al. 2006). Data reduction was carried out using a modified *Herschel* interactive processing environment (HIPE) 3.0 pipeline, starting from the level 0 data and produced maps with pixel sizes of 3.2, 3.2 and 6.4'' at 70, 100 and 160 μm with PSF FWHM values of 5.2, 7.7 and 12'' respectively. HIPE is used to suppress the bad pixels and those affected by saturation and convert the signal to Jy pixel⁻¹. We perform flatfield correction and apply astrometry to the data

* *Herschel* is an ESA space observatory with science instruments provided by Principal Investigator consortia. It is open for proposals for observing time from the worldwide astronomical community.

** Figure 5 is only available in electronic form at <http://www.aanda.org>.

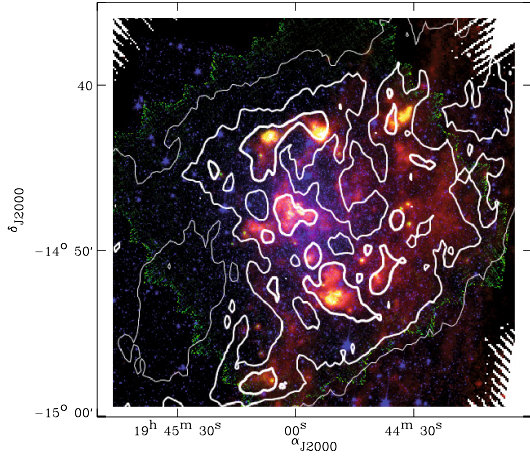


Fig. 1. NGC 6822 *Herschel/Spitzer* 3-color image. North is up, East is left. *Blue*: stellar emission observed with *Spitzer/IRAC* 3.6 μm . *Yellow*: warm dust emission with *Herschel/PACS* 100 μm . *Red*: cold dust emission with *Herschel/SPIRE* 250 μm . H I contours at 3.1, 7.8 and 13×10^{20} H cm^{-2} are overlaid (de Blok & Walter 2000). The two bright compact IR knots in the North are Hubble X (east) and Hubble V (west).

cube. We then apply a multiresolution median transform (MMT) deglitching correction and a second order deglitching process to the data. We perform polynomial fits on half scan legs to subtract the baselines. The galaxy was masked in the data cube during this step to avoid an overestimation of the signal level on the source. The median baseline subtraction step suppresses most of the bolometer temperature drifts. We generate the final maps using the madmap method of HIPE. The absolute flux calibration uncertainties are estimated to be $\pm 10\%$.

SPIRE (Griffin et al. 2010) observations, at 250, 350 and 500 μm , cover a region of $26' \times 26'$. Data have been reprocessed from the level 1 cube following the steps described in Pohlen et al. (2010) and Bendo et al. (2010). The overall absolute calibration accuracy is estimated to be $\pm 15\%$ (Swinyard et al. 2010). The pipeline produces maps with a pixel size of 6'', 10'' and 14'' at 250, 350 and 500 μm with PSF *FWHM* values of 18'', 25'' and 36'' respectively. The SPIRE ICC has released preliminary multiplying factors to correct for extended sources: 1.02, 1.05 and 0.94 at 250, 350 and 500 μm ¹.

3. Analysis

3.1. *Herschel* maps

Figure 1 is a color composition (size of the SPIRE maps) showing how stars, dust and H I gas (in contours) are distributed in the galaxy. IRAC 3.6 μm band (blue) mostly traces the stellar emission while PACS 100 μm (green) traces the warm dust. SPIRE 250 μm (red) traces the cold dust phases. We find that the bright H II regions are resolved at all *Herschel* wavelengths. The brightest IR/submm SF regions all coincide with H I peaks. While H I extends 30' to the NW and SE of the galaxy, the big H I hole is devoid of dust emission. The star formation history of NGC 6822 began 12–15 Gyr ago and has been quiescent until about 0.6–1 Gyr ago (Gallart et al. 1996; Wyder 2001). Bianchi et al. (2001) found very young stellar populations (<10 Myr) in the star-forming regions Hubble V and X. These regions are also

¹ As advised by the ICC for SD papers, see http://herschel.esac.esa.int/SDP_wkshops/presentations/IR/3_Griffin_SPIRE_SDP2009.pdf

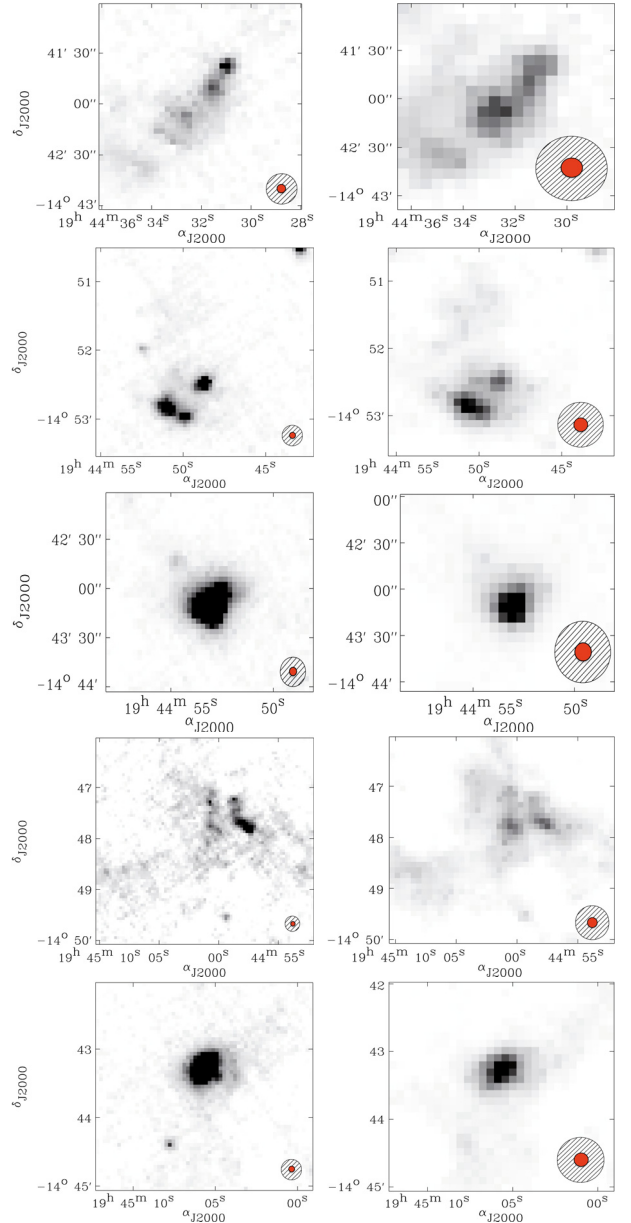


Fig. 2. PACS 70 μm (left) and 160 μm (right) observations of bright star-forming regions of NGC 6822. These regions are numbered 3, 7, 8 (Hubble V), 11 and 14 (Hubble X) in Fig. 3a and in this paper. The *FWHM* of *Spitzer/MIPS* 70 and 160 μm PSFs (black striped circles) and those of PACS 70 and 160 μm (red circles) are overlaid for comparison.

the brightest structures of the FIR and submm emission (see the two brightest knots in the North of NGC 6822 in Fig. 1).

PACS 70 and 160 μm observations of the 5 brightest star-forming regions of NGC 6822 are shown in Fig. 2. These regions are respectively numbered 3, 7, 8, 11 and 14 in Fig. 3a. The increase in spatial resolution from MIPS (Cannon et al. 2006) to PACS (see the respective *FWHM* PSFs in Fig. 2) now enables us to nicely separate the different substructures of the star-forming regions 3 and 11 (see online material for complete images) but also shows faint emission across the galaxy that was detected but not resolved with MIPS. To compare MIPS and PACS fluxes at 70 and 160 μm for these H II regions, we convolve the images to the lowest (*FWHM* MIPS 160 μm : 40'') resolution (Bendo et al. 2010) and extract flux densities using the function *aper* of IDL in apertures of 57'' radius (~ 135 pc). We find that, for the bright star-forming regions of NGC 6822, MIPS and PACS flux

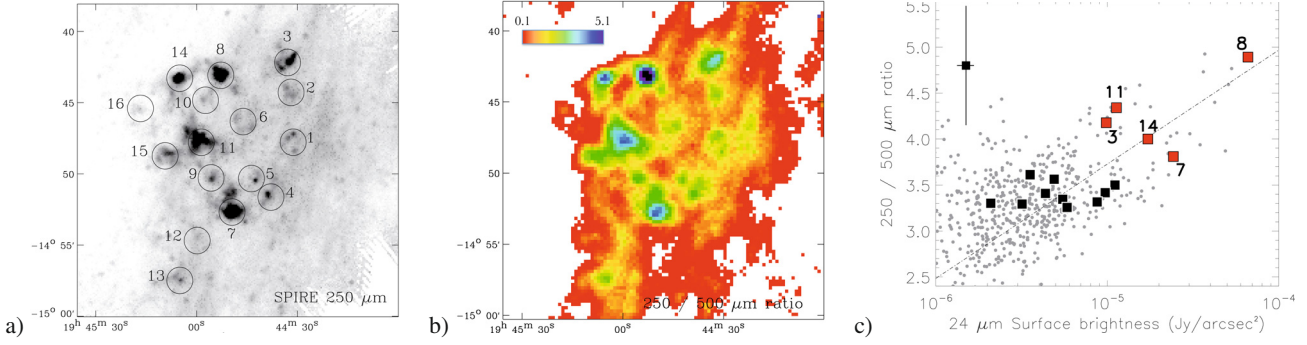


Fig. 3. **a)** NGC 6822 observed at 250 μm . Circles indicate our photometric apertures (57'' radius). **b)** 250/500 μm flux density ratio map of NGC 6822. **c)** $F_{\nu}(250)/F_{\nu}(500)$ ratio as a function of the 24 μm surface brightness. Ratios are estimated in the apertures of Fig. 3a (squares) and pixel-by-pixel (grey circles). In this second case, pixels are chosen to have the size of the *FWHM* of SPIRE 500 μm . Red squares indicate the brightest H II regions. Since the uncertainties in the SPIRE fluxes are preliminary, we give an indication of the error bars on the upper left corner.

densities compare within $\pm 25\%$ at 70 and 160 μm but 33% for Hubble V at 70 μm . The SPIRE maps also resolve the structures detected with PACS. We find that the diffuse emission of NGC 6822 is strongly affected by non homogeneous Galactic cirrus emission due to the low Galactic latitude of NGC 6822 (low level diffuse emission in Fig. 3a). We model the cirrus emission in each map as a plane after first masking the emission that we associate with the galaxy. We then remove the modeled cirrus emission from each SPIRE map for the analysis. The cirrus contamination is estimated to be less than 10% of the flux densities of the star-forming regions at 250 and 350 μm . It is less than 15% and 25% on the less active regions respectively at 250 μm and 350 μm . The contamination is finally estimated to be $\sim 25\%$ on the bright regions at 500 μm but can be as high as $\sim 50\%$ of the emission in the lowest level of the diffuse ISM. Our quantitative results concern relatively bright regions. Thus, taking into account the various uncertainties, we attribute an overall conservative estimate of $\sim 30\%$ for all SPIRE bands. We note that our results are strongly dependant on the SPIRE fluxes, more precisely on our treatment of cirrus emission that could have been underestimated if some of this emission is more significant along some lines of sight.

3.2. SPIRE band ratios

We convolve the 250 μm map to the resolution of SPIRE 500 μm (36'') and build a 250/500 flux density ratio map (Fig. 3b) to study the evolution of the submm regime of the SEDs. The ratio peaks in the H II regions (Hubble V showing the highest ratio). This ratio map highlights the evolution of the dust temperature distribution across the galaxy: warmer toward the H II regions and decreasing in the diffuse extended dust component between the H II regions. To understand the processes contributing to the heating of dust, we examine how the submm part of the SED evolves with star formation. We select individual regions and estimate the SPIRE 250 and 500 μm flux densities of these regions in apertures of 57'' radius (Fig. 3a), corresponding to regions studied in Cannon et al. (2006). In Fig. 3c, we plot the 250/500 flux density ratio of the selected regions (squares) as a function of their 24 μm surface brightness. The same 250/500 ratios performed on a pixel-by-pixel² basis throughout the whole map are overlaid (grey circles). The 24 μm flux is commonly used as a tracer of star formation (e.g. Calzetti 2007). The 250/500 ratios seem to correlate with the

24 μm surface brightness across NGC 6822, which could imply that the cold dust temperature distribution varies with the star formation activity of the region, with higher temperature dust present where star formation activity dominates (Boselli et al. 2010). Bendo et al. (2010) find that the SPIRE band ratios in M81 depend on radius, and that the old stellar population of the bulge and disk could be the primary source for the dust emission seen by SPIRE. Their submm ratios do not show any strong correlation with 24 μm surface brightness. This effect suggests that the dust heating processes in low metallicity starbursts may differ from normal dusty spirals. The ISM of dwarfs is, indeed, less opaque than spirals, with dense SF regions usually influencing the whole galaxy. These galaxies also tend to have a preponderance of younger stellar population and less evolved stars, and thus very different star formation histories.

3.3. SEDs

To study the variations of the local SEDs, we select two H II regions (Hubble V and X, numbered 8 and 14 in Fig. 3a) and three less active regions (Reg. 6, 9 and 13). We suspect the median baseline filtering of the PACS reduction to remove some of the diffuse emission. PACS fluxes are thus not used in the modelling of the total SED and less active regions. *Spitzer* observations (SINGS 5th release, Kennicutt et al. 2003) complete the coverage. We convolve the images to MIPS 160 μm resolution (40'') and estimate the flux densities within 57'' radius apertures.

We use a realistic SED model which follows the approach of the Dale et al. (2001) and Draine & Li (2007) models, and use the dust composition and size distribution of Zubko et al. (2004). The interstellar radiation field is assumed to have the shape of the Galactic diffuse ISM of Mathis et al. (1983). The dust mass exposed to a given heating intensity U is given by: $dM_{\text{dust}}(U) \propto U^{-\alpha} dU$ with $U_{\text{min}} < U < U_{\text{max}}$ (Dale et al. 2001). α parametrises the contribution of the different local SEDs exposed to U . Details on the modelling can be found in Galametz et al. (2009). Serra Díaz-Cano & Jones (2008) study carbon dust in shock waves and warn about using graphites in dust models. To study how this choice affects our dust masses, we test both graphites and amorphous carbons (Rouleau & Martin 1991) to describe the interstellar carbon dust. Figure 4 presents the global SED of NGC 6822 along with the individual SEDs of the 5 selected regions obtained with amorphous carbons. No submm excess seems to be detected in NGC 6822, contrary to other dwarf galaxies observed with *Herschel* (O'Halloran et al. 2010; Grossi et al. 2010). We find that the SEDs of H II regions have warmer dust temperatures than less active regions. The total SED

² The pixel size of the maps was chosen to equal the *FWHM* of SPIRE 500 μm (36'').

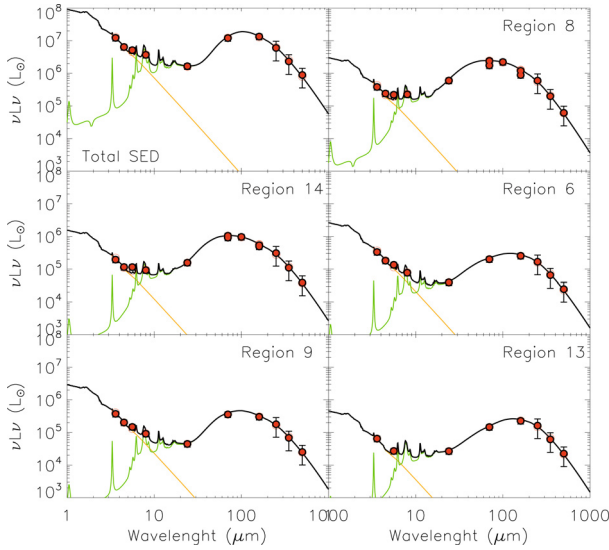


Fig. 4. Total SED of NGC 6822 and SEDs of individual regions among which bright H II regions (Region 8 and 14) and less active regions (Regions 6, 9 and 13). Our SED models are plotted in black. Observational constraints are overlaid in red circles. The orange and green lines distinguish the stellar and the dust contribution. The 30% uncertainties are conservatively estimated for the SPIRE bands.

Table 1. General properties.

| Id | M_{HI} [$10^5 M_{\odot}$] | $M_{\text{dust Gr}}^a$ [$10^3 M_{\odot}$] | $M_{\text{HI}}/M_{\text{dust Gr}}$ | $M_{\text{dust AC}}^a$ [$10^3 M_{\odot}$] | $M_{\text{HI}}/M_{\text{dust AC}}$ |
|-------|---|--|------------------------------------|--|------------------------------------|
| 8 | 7.6 | 22.2^{+7}_{-6} | 34 | 7.8^{+3}_{-3} | 97 |
| 14 | 9.1 | 23.0^{+8}_{-7} | 39 | 10.2^{+4}_{-4} | 89 |
| 6 | 5.8 | 14.1^{+5}_{-4} | 41 | 5.3^{+2}_{-2} | 109 |
| 9 | 7.7 | 22.1^{+5}_{-5} | 34 | 8.5^{+3}_{-4} | 91 |
| 13 | 7.8 | 12.9^{+5}_{-3} | 60 | 4.7^{+3}_{-3} | 166 |
| Total | 500 ^b | 612^{+105}_{-182} | 80 | 269^{+143}_{-145} | 186 |

Notes. ^(a) Dust mass derived using graphite (Gr) or amorphous carbon (AC). ^(b) H I mass corresponding to the region mapped with *Herschel*.

of NGC 6822 does not show a high 24/70 ratio, indicating that it may not be dominated by the IR emission of bright H II regions. The dust masses derived using graphites are 2.2 to 2.8 times higher than those using amorphous carbons (Table 1), due to their lower emissivity at submm wavelength.

To estimate the gas-to-dust mass ratios (G/D), we derive the H I mass of our regions from the integrated map of de Blok & Walter (2000). We also estimate the H I mass corresponding to the region mapped with *Herschel* to be $\sim 5 \times 10^7 M_{\odot}$. Gratier et al. (2010) found that H₂ derived from CO observations should not represent more than $10^7 M_{\odot}$. Faint emission lines from warm H₂ are observed in Hubble V (Hunter & Kaufman 2007). Cannon et al. (2006) note that the major H II regions correspond to the strongest H α sources of the galaxy. From their H α fluxes, we derive an H α mass inferior to $10^5 M_{\odot}$ in the H II regions (Storey & Hummer 1995, assuming $T = 10^4$ K and $N_e = 100$). H I thus dominates the gas mass in NGC 6822. The Galliano et al. (2008) models predicts G/D of ~ 500 for galaxies presenting the metallicity of NGC 6822. We find low G/D (Table 1) for the individual regions compared to what can be expected from

chemical evolution models, especially when graphite grains are used in the modelling. Amorphous carbon results in a flatter submm slope compared to graphite and thus requires less mass to produce the same emission. These results are consistent with those found by Meixner et al. (2010) in the Large Magellanic Cloud. The total dust mass of the central region mapped with *Herschel* is $2.7 \times 10^5 M_{\odot}$ using amorphous carbon dust, leading to a total G/D of 186.

4. Conclusions

We present *Herschel* images of NGC 6822 which resolve ISM structures up to $500 \mu\text{m}$. We find that the 250/500 ratio (tracing the cold dust temperature range) may be dependent on the $24 \mu\text{m}$ surface brightness and thus trace the SF activity. We model individual SEDs across NGC 6822 and show that the SED shape is evolving from H II regions to less active regions, with H II regions having a warmer dust temperature range. We derive very high dust masses using graphite to describe carbon dust and find that the use of amorphous carbon decreases the dust masses, indicating that SED models including *Herschel* constraints require different dust properties, namely more emissive grains.

Acknowledgements. We thank the referee for his comments that help to improve the quality of this paper. We also thank Erwin de Blok for the integrated H I map of NGC 6822. PACS has been developed by MPE (Germany); UVIE (Austria); KU Leuven, CSL, IMEC (Belgium); CEA, LAM (France); MPIA (Germany); INAF-IFSI/OAA/OAP/OAT, LENS, SISSA (Italy); IAC (Spain). This development has been supported by BMVIT (Austria), ESA-PRODEX (Belgium), CEA/CNES (France), DLR (Germany), ASI/INAF (Italy), and CICYT/MCYT (Spain). SPIRE has been developed by Cardiff University (UK); Univ. Lethbridge (Canada); NAOAC (China); CEA, LAM (France); IFSI, Univ. Padua (Italy); IAC (Spain); SNSB (Sweden); Imperial College London, RAL, UCL-MSSL, UKATC, Univ. Sussex (UK) and Caltech, JPL, NHSC, Univ. Colorado (USA). This development has been supported by CSA (Canada); NAOAC (China); CEA, CNES, CNRS (France); ASI (Italy); MCINN (Spain); Stockholm Observatory (Sweden); STFC (UK); and NASA (USA).

References

- Bendo, G. J., et al. 2010, A&A, 518, L65
 Bianchi, L., Scuderi, S., Massey, P., & Romaniello, M. 2001, AJ, 121, 2020
 Boselli, A., et al. 2010, A&A, 518, L61
 Calzetti, D. 2007, Nuovo Cimento B Ser., 122, 971
 Cannon, J. M., Walter, F., Armus, L., et al. 2006, ApJ, 652, 1170
 Dale, D. A., Helou, G., Contursi, A., Silberman, N. A., & Kolhatkar, S. 2001, ApJ, 549, 215
 de Blok, W. J. G., & Walter, F. 2000, ApJ, 537, L95
 Draine, B. T., & Li, A. 2007, ApJ, 657, 810
 Galametz, M., Madden, S., Galliano, F., et al. 2009, A&A, 508, 645
 Gallart, C., Aparicio, A., Bertelli, G., & Chiosi, C. 1996, AJ, 112, 1950
 Galliano, F., Dwek, E., & Chianial, P. 2008, ApJ, 672, 214
 Gratier, P., Braine, J., Rodriguez-Fernandez, N. J., et al. 2010, A&A, 512, A68
 Griffin, M. J., et al. 2010, A&A, 518, L3
 Grossi, M., et al. 2010, A&A, 518, L52
 Hunter, D. A., & Kaufman, M. 2007, AJ, 134, 721
 Kennicutt, Jr., R. C., Armus, L., Bendo, G., et al. 2003, PASP, 115, 928
 Mateo, M. L. 1998, ARA&A, 36, 435
 Mathis, J. S., Mezger, P. G., & Panagia, N. 1983, A&A, 128, 212
 Meixner, M., et al. 2010, A&A, 518, L71
 O'Halloran, B., et al. 2010, A&A, 518, L58
 Pilbratt, G. L., et al. 2010, A&A, 518, L1
 Poglitsch, A., et al. 2010, A&A, 518, L2
 Pohlen, M., et al. 2010, A&A, 518, L72
 Rouleau, F., & Martin, P. G. 1991, ApJ, 377, 526
 Serra Díaz-Cano, L., & Jones, A. P. 2008, A&A, 492, 127
 Storey, P. J., & Hummer, D. G. 1995, MNRAS, 272, 41
 Swinyard, B. M., et al. 2010, A&A, 518, L4
 Wyder, T. K. 2001, AJ, 122, 2490
 Zubko, V., Dwek, E., & Arendt, R. G. 2004, ApJS, 152, 211

-
- ¹ CEA, Laboratoire AIM, Irfu/SAP, Orme des Merisiers, 91191 Gif-sur-Yvette, France
e-mail: maud.galametz@cea.fr
- ² School of Physics & Astronomy, Cardiff University, Queens Buildings The Parade, Cardiff CF24 3AA, UK
- ³ Astrophysics Group, Imperial College, Blackett Laboratory, Prince Consort Road, London SW7 2AZ, UK
- ⁴ Sterrenkundig Observatorium, Universiteit Gent, Krijgslaan 281 S9, 9000 Gent, Belgium
- ⁵ Dept. of Physics & Astronomy, University College London, Gower Street, London WC1E 6BT, UK
- ⁶ Jet Propulsion Laboratory, CA 91109, Dept. of Astronomy, California Institute of Technology, CA 91125 Pasadena, USA
- ⁷ Laboratoire d'Astrophysique de Marseille, UMR6110 CNRS, 38 rue F. Joliot-Curie, 13388 Marseille, France
- ⁸ Instituto de Astrofísica de Canarias (IAC) and Dept. de Astrofísica, Universidad de La Laguna (ULL), La Laguna, Tenerife, Spain
- ⁹ Institut d'Astrophysique de Paris, UMR7095 CNRS, Univ. Pierre & Marie Curie, Boulevard Arago, 75014 Paris, France
- ¹⁰ Dept. of Physics & Astronomy, University of California, Irvine, CA 92697, USA
- ¹¹ Observational Cosmology Lab, Code 665, NASA Goddard Space Flight Center Greenbelt, MD 20771, USA
- ¹² Dept. of Astrophysical & Planetary Sciences, CASA CB-389, Univ. of Colorado, Boulder, CO 80309, USA
- ¹³ ESA Astrophysics Missions Division, ESTEC, PO Box 299, 2200 AG Noordwijk, The Netherlands
- ¹⁴ Astronomy Centre, Department of Physics & Astronomy, Univ. of Sussex, UK
- ¹⁵ Mullard Space Science Laboratory, University College London, Holmbury St Mary, Dorking, Surrey RH5 6NT, UK
- ¹⁶ Dept. of Physics & Astronomy, McMaster University, Hamilton, Ontario, L8S 4M1, Canada
- ¹⁷ School of Physics & Astronomy, Univ. of Nottingham, University Park, Nottingham NG7 2RD, UK
- ¹⁸ Istituto di Fisica dello Spazio Interplanetario, INAF, Via del Fosso del Cavaliere 100, 00133 Roma, Italy
- ¹⁹ Infrared Processing & Analysis Center, California Institute of Technology, Mail Code 100-22, 770 South Wilson Av, Pasadena, CA 91125, USA
- ²⁰ Centre for Astrophysics Research, Univ. of Hertfordshire, College Lane, Herts AL10 9AB, UK
- ²¹ University of Padova, Department of Astronomy, Vicolo Osservatorio 3, 35122 Padova, Italy
- ²² Observatoire Astronomique de Strasbourg, UMR 7550 Univ. de Strasbourg – CNRS, 11, rue de l'Université, 67000 Strasbourg, France
- ²³ UK Astronomy Technology Center, Royal Observatory Edinburgh, Edinburgh, EH9 3HJ, UK
- ²⁴ Institut für Astronomie, Universität Wien, Trkenschanzstr. 17, 1180 Wien, Austria

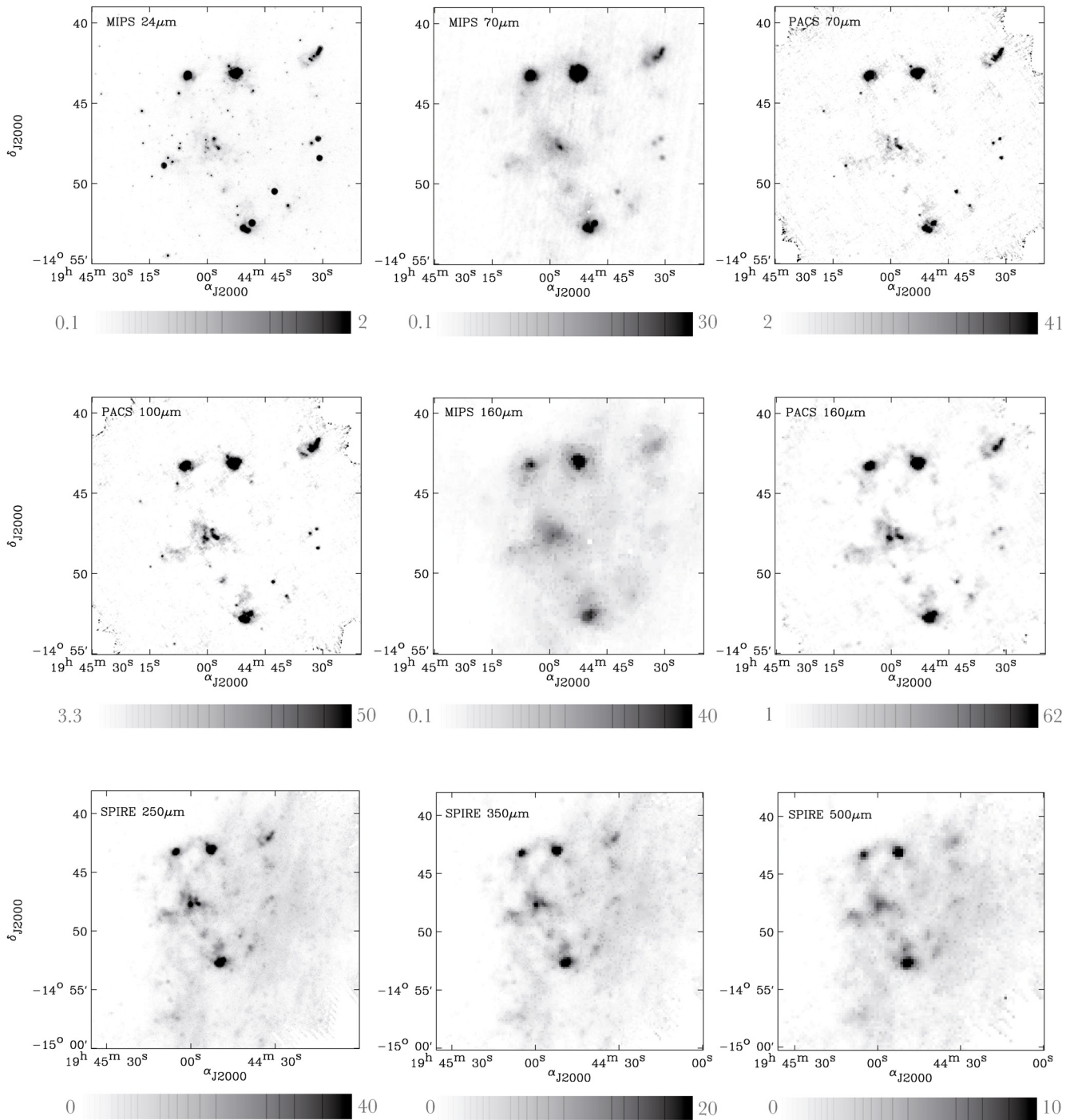


Fig. 5. NGC 6822 observed by *Spitzer*/MIPS and *Herschel*/PACS and SPIRE. Fluxes are in MJy/sr.